Table 1. List of 36 (candidate) UCXBs (adapted from Nelemans & Jonker 2006a and in 't Zand et al. 2007), including 7 cases proposed on the basis of very low $L_{\rm X}$. We leave out cases identified through the diagnostic of the X-ray continuum parameter values (Sidoli et al. 2001), for instance EXO 1745-248 in Terzan 5 (Heinke et al. 2003), because that diagnostic is not always consistent with others (e.g., Verbunt & Lewin 2006).

Name	(1)	(2)	(3)	$P_{\rm orb}$ (min)
13 certain UCXBs				
MAXI J0911-655 (in NGC 2808)	pp	T	P	44[aa]
XTE J0929-314	pp	T	M	44[a]
4U 1626-67	pp	P	P	42[b]
2FGL J1653.6-0159	po	?	?	75[y]
IGR J16597-3704	pp	T	M	46[w]
IGR J17062-6143	pp	P	B	40[x]
NGC 6440 X-2	pp	T	M	57[z]
XTE J1751-305	pp	T	M	42[c]
SWIFT J1756.9-2508	pp	T	M	55[u]
XTE J1807-294	pp	T	M	40[d]
4U 1820-303 (in NGC 6624)	px	P	B	11[e]
4U 1850-087 (in NGC 6712)	po	P	B	21 or 13[f]
4U 1915-05	px	P	B,D	50[g]
M15 X-2 (in M15)	po	P	B	23[h]
7 candidate UCXBs with tentative orbital periods				
4U 0513-40 (in NGC 1851)	po,r[l]	P	B	17[r]
4U 0614+091	po,r	P	B	50[i]
2S 0918-549	po,r[l]	P	B	17[q]
4U 1543-624	po	P	B	18[j]
4U 1728-34	px	P	B	11[t]
H 1825-331 (in NGC 6652)	po	P	B	55[k]
NGC 6652 B (in NGC 6652)	po	Q		44[k]
5 candidate UCXBs with low optical to X-ray flux				
1A 1246-588	r[m],x	P	B	
4U 1812-12	r[m],x	P	B	
4U 1822-000	r[l]	P	B	
4U 1905+000	r[n]	T	B	
ω Cen qLMXB	r[o]	Q		
9 candidate UCXBs based on persistence & low $M-dot$				
SAX J1712.6-3739	x	P	B	
1RXS J170854.4-321857	x[p]	P	B	
1RXS J171824.2-402934	x[p]	P	B	
4U 1722-30 (in Terzan 2)	x	P	B	
1RXS J172525.2-325717	x	P	B	
SLX 1735-269	x	P	B	
SLX 1737-282	x	P	B	
SLX 1744-299	x	P	B	
XMMU J174716.1-281048	x[s]	LT	B	
1 candidate UCXB based on 0.6 keV emission line				
RX J170930.2-263927	x[v]	T	B	

(1) Type of argument for ultracompact identification: r = $L_{\rm x}/L_{\rm opt}$ argument, p = period measurement (pp=pulsar, px=dips/eclipse, po=optical modulation), x = persistent burster with low L; (2) Type of accretion: P = persistent, T = transient, Q = quiescent (never seen to be luminous), LT=long transient; (3) Type of source: P = pulsar, M = accretion-powered millisecond pulsar, B = burster, D = eclipser and/or dipper; [a]Galloway et al. (2002); [b]Middleditch et al. 1981; [c]Markwardt et al. (2002); [d]Markwardt et al. 2003; [e]Stella et al. 1987; [f]Homer et al. 1996; [g]White & Swank 1982; [h]Dieball et al. 2005; [i]Shahbaz et al. 2008; [j]Wang & Chakrabarty (2004); [k]Heinke et al. 2001; [l]Juett et al. 2001; [m]Bassa et al. 2006; [n]Jonker et al. 2006; [o]Haggard et al. 2004; [p]in 't Zand et al. 2005a; [q]Zhong & Wang 2011; [r]Zurek et al. 2009; [s]Degenaar et al. 2011; [t]Galloway et al. 2010; [u]Krim et al. 2007; [v]Jonker et al. 2003; [w]Sanna et al. 2018; [x]Strohmayer et al. 2018; [y]Kong et al. 2014; [z]Altamirano et al. 2010
[aa]Sanna et al. 2017

Compiled by Jean in 't Zand, dd. May 23, 2020.